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TriAnd and its siblings: satellites of satellites in the Milky Way halo

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ABSTRACT

We explore the Triangulum–Andromeda (TriAnd) overdensity in the SPLASH (Spectroscopic and Photometric Landscape of Andromeda’s Stellar Halo) and SEGUE (the Sloan Extension for Galactic Understanding and Exploration) spectroscopic surveys. Milky Way main-sequence turn-off stars in the SPLASH survey reveal that the TriAnd overdensity and the recently discovered Pan-Andromeda Archaeological Survey (PAndAS) stream share a common heliocentric distance ($D \sim 20$ kpc), position on the sky, and line-of-sight velocity ($V_{\text{GSR}} \sim 50 \text{ km s}^{-1}$). Similarly, A-type, giant, and main-sequence turn-off stars selected from the SEGUE survey in the vicinity of the *Segue* 2 satellite show that TriAnd is prevalent in these fields, with a velocity and distance similar to *Segue* 2. The coincidence of the PAndAS stream and *Segue* 2 satellite in positional and velocity space to TriAnd suggests that these substructures are all associated, and may be a fossil record of group-infall on to the Milky Way halo. In this scenario, the *Segue* 2 satellite and PAndAS stream are ‘satellites of satellites’, and the large, metal-rich TriAnd overdensity is the remains of the group central.

Key words: Galaxy: formation – Galaxy: halo – galaxies: dwarf.

1 INTRODUCTION

In the modern cold dark matter cosmology, dark matter haloes form hierarchically, and the presence of ‘subhaloes’ or ‘sub-subhaloes’ is naturally predicted. Dark matter subhaloes are likened to the satellite galaxies that orbit the Milky Way (MW) galaxy, and the agreement/disagreement between the observed satellite population and the predictions from simulations are well documented (e.g. Flores & Primack 1994; Moore 1994; Klypin et al. 1999; Moore et al. 1999). However, the population of sub-subhaloes (or satellites of satellites) is less well studied. In essence, smaller constituents should be associated with the satellite population today and/or the destroyed remnants of accreted satellites in the stellar halo. It remains to be seen what fraction of these sub-subhaloes have survived to the present day, or even if we are able to observe them (i.e. if they are too faint, or completely ‘dark’).

Observational evidence for associations between satellites and/or streams increased dramatically following the discovery of very low luminosity galaxies ($L \lesssim 10^5 L_{\odot}$) in the MW (e.g. Willman et al. 2005; Belokurov et al. 2006b, 2007a). Amongst the faintest of these, curious associations in both position and velocity space with other satellites or streams have been found. For example, Newberg et al. (2010) showed that the Orphan stellar stream (Belokurov et al. 2006a; Grillmair 2006; Belokurov et al. 2007c) has a similar distance and velocity as the *Segue* 1 satellite, and is only ~ 2 deg away on the sky. The narrow width of the Orphan stream (~ 1 deg) indicates that *Segue* 1 is probably not the progenitor of the stream, but it is likely that the two are dynamically associated. The proximity of the Boötes II satellite, in both position and velocity space, with the Sagittarius stream led Koch et al. (2009) to suggest that this ultrafaint may have been stripped from the more massive Sagittarius dwarf. Similarly, Belokurov et al. (2009) noted the nearness of the *Segue* 2 satellite (perhaps the ‘least massive galaxy’ known; Kirby et al. 2013a) to the elusive Triangulum–Andromeda (TriAnd) overdensity (see below), which again suggests that a very faint satellite may be associated with a more massive structure. Finally, the two ultrafaints, Leo IV and Leo V, discovered in quick succession by

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Belokurov et al. (2007a, 2008) are separated by ~ 3 deg on the sky, with small offsets in both distance and velocity (~ 20 kpc, ~ 40 km s $^{-1}$); Belokurov et al. (2008) dub Leo V as ‘a companion of a companion of the MW galaxy’.

The enhanced evidence for satellite–satellite associations from the lowest mass galaxies is perhaps unsurprising. Wetzel, Deason & Garrison-Kimmel (in preparation) recently showed that a significant fraction (~ 30 per cent) of low-mass subhaloes ($M_{\text{star}} \lesssim 10^5 M_{\odot}$) likely fell into a MW-type host as a satellite of a more massive subhalo. These results suggest that the most likely place to find associations between satellites and streams today is in the vicinity of the most massive substructures in the Galaxy, which may have hosted less massive sub-subhaloes at the time of infall.

Among the various streams/overdensities identified in the MW halo, there are five ‘big ones’ that stand out the most: the Sagittarius stream (Sgr; Ibata, Gilmore & Irwin 1994; Yanny et al. 2000; Majewski et al. 2003), the Monoceros Ring¹ (Newberg et al. 2002; Ibata et al. 2003), the Hercules Aquila cloud (Belokurov et al. 2007b; Simion et al. 2014), the Virgo overdensity (Vivas 2002; Jurić et al. 2008; Bonaca et al. 2012a), and the TriAnd stream (Majewski et al. 2004; Rocha-Pinto et al. 2004; Sheffield et al. 2014). It is likely that these are the remnants of fairly massive accretion events, and they each span several hundreds of degrees on the sky. Most of these structures have been studied extensively, largely thanks to the wide sky coverage of the Sloan Digital Sky Survey (SDSS). However, TriAnd, which is the least well known, and perhaps the smallest of the five, has evaded significant scrutiny in the literature.

The TriAnd overdensity, so-called from its discovery in the constellations of TriAnd (Majewski et al. 2004; Rocha-Pinto et al. 2004), is a large, diffuse structure that lies at a heliocentric distance of $10 \lesssim D/\text{kpc} \lesssim 30$. The full extent of this overdensity is poorly known, it subtends at least $50^{\circ} \times 60^{\circ}$ on the sky, but is located in a region of the sky that has not been completely covered by large surveys such as SDSS. As a result, the TriAnd overdensity has received relatively little attention. In fact, most of the work on TriAnd has resulted from serendipitous observations focused on the M31 galaxy (Majewski et al. 2004; Martin, Ibata & Irwin 2007; Martin et al. 2014). The nature of TriAnd has been debated in the literature; Rocha-Pinto et al. (2004) propose that it is either a bound core of a very dark-matter dominated dwarf or a portion of a tidal stream, while Johnston et al. (2012) suggest that TriAnd is similar to ‘cloud-like’ debris from a disrupted dwarf galaxy (see also Deason et al. 2013; Sheffield et al. 2014).

TriAnd is not the only MW substructure discovered from M31 surveys; Martin et al. (2014) recently discovered a new faint stellar stream located at $D \sim 20$ kpc in their photometric survey of M31. Most surveys of the MW halo tend to avoid regions close to the Galactic disc, so this detection suggests there may be several low latitude streams awaiting discovery. The stream is dynamically cold ($\sigma < 7.1$ km s $^{-1}$) and has a width of 300–650 pc, which lead Martin et al. (2014) to argue for a dwarf-galaxy-accretion origin.

In this contribution, we study the TriAnd overdensity in two large spectroscopic surveys, Spectroscopic and Photometric Landscape of Andromeda’s Stellar Halo (SPLASH) and Sloan Extension for Galactic Understanding and Exploration (SEGUE), and argue for the association of a stream [the newly discovered Pan-Andromeda Archaeological Survey (PAndAS) stream] and a satellite (*Segue 2*)

with this more massive structure. The paper is arranged as follows. Section 2 briefly describes the spectroscopic surveys used in this work. In Sections 3 and 4, we use SPLASH and SEGUE to identify stars likely belonging to the TriAnd overdensity along the line of sight towards M31, and in the vicinity of the *Segue 2* satellite. We discuss the implications of our results in Section 5, and we summarize our main conclusions in Section 6.

2 SPECTROSCOPIC DATA

2.1 Spectroscopic and Photometric Landscape of Andromeda’s Stellar Halo (SPLASH)

The SPLASH data that we use in this work are identical to the data used by Gilbert et al. (2012) to characterize the M31 stellar halo. Here, we provide a brief summary of the observations but refer the interested reader to earlier SPLASH papers for more details (Guhathakurta et al. 2006; Gilbert et al. 2012).

The photometric catalogues used to target M31 red giant branch (RGB) stars for spectroscopic follow-up primarily came from imaging observations obtained with the Mosaic Camera on the Kitt Peak National Observatory (KPNO) 4-m Mayall telescope. Observations were made in the Washington system (M , T_2 , $DD051$), which can be used to efficiently separate giant stars in M31 from foreground MW stars (Majewski et al. 2000). Additional photometry was compiled using V and I (or g' and i') filters obtained from observations with the William Herschel Telescope, the Suprime-Cam instrument on the Subaru Telescope, and the MegaCam instrument on the 3.6-m Canada–France–Hawaii Telescope (CFHT).

The M31 fields span a large range in position angle and projected distance from the centre of M31 (see Fig. 1). The above photometric catalogues were used to obtain spectroscopic observations

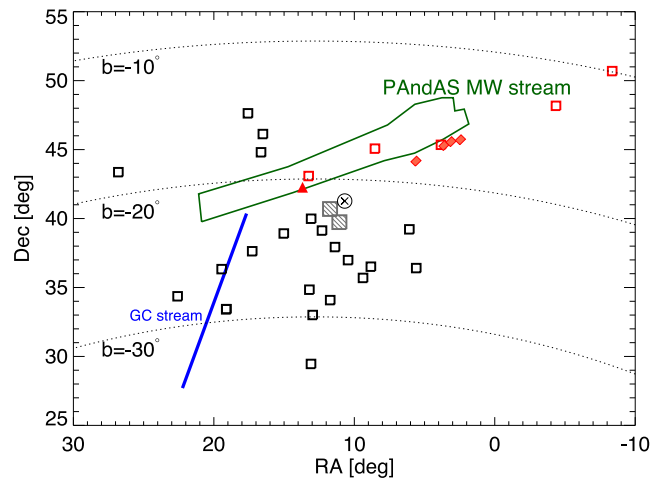


Figure 1. Locations of SPLASH fields with spectroscopic measurements included in Gilbert et al. (2012). We show the footprints of the KPNO/Mosaic (open squares), CFHT/MegaCam (line-filled grey squares), and Subaru/Suprime-Cam (tilted filled rectangles) images used to target stars for spectroscopic follow-up. The location of the PAndAS stream is indicated with the green region, and fields coincident with the PAndAS stream are highlighted in red. The red triangle indicates the field where Martin et al. (2014) measured radial velocities for the PAndAS stream stars. The centre of M31 is indicated with a cross symbol, and the dotted lines show lines of constant Galactic latitude. The location of the recently discovered globular cluster stream by Bonaca, Geha & Kallivayalil (2012b) is shown with the blue line. The TriAnd overdensity likely fills most of the region shown (cf. fig. 2 in Martin et al. 2014).

¹ Note that it has also been claimed that the Monoceros Ring is the signature of disc flaring rather than the signature of a disrupted dwarf galaxy (e.g. Ibata et al. 2003; Momany et al. 2006).

of M31 stars using the DEep Imaging Multi-Object Spectrograph (DEIMOS) on the Keck II telescope. M31 stars were given higher priority for inclusion on the slit mask, but lower priority filler targets, such as MW foreground stars, were often included on the masks due to the paucity of M31 RGB stars (especially in the low density, outer fields of M31). Spectra were obtained over nine observing seasons (Fall 2002–2010), with typical integration times of 1 h per mask. The 1200 line mm^{-1} grating was used for all observations, and the typical wavelength range of the spectra is 6450–9150 Å.

The reduction of the spectra using the SPEC2D and SPEC1D software (Cooper et al. 2012; Newman et al. 2013) is described in more detail in Gilbert et al. (2012). The resulting reduced spectra from 108 masks yielded successful velocity measurements for over 5800 stellar spectra. In this work, we are only interested in the MW foreground spectra, specifically the turn-off stars in the MW halo (see Section 3).

2.2 Sloan Extension for Galactic Understanding and Exploration (SEGUE)

The SDSS (York et al. 2000) is an imaging and spectroscopic survey covering over one quarter of the sky. A dedicated 2.5 m telescope at Apache Point Observatory, New Mexico is used to obtain *ugriz* imaging (Fukugita et al. 1996) and moderate-resolution ($R \sim 2000$) spectra. SEGUE (including both SEGUE 1 and 2), a key project executed during SDSS-II and SDSS-III, has obtained spectra for $\sim 360\,000$ stars in the MW, and was designed to explore the kinematics and populations of our Galaxy and its halo (Yanny et al. 2009).

The wavelength coverage of the SEGUE survey spans 3900–9000 Å and targets fainter MW stars ($14.0 < g < 20.4$) of a wide variety of spectral types. The stellar populations include main-sequence stars, hot A-type stars, and evolved giant stars. The SEGUE spectra are clustered in ‘pencil-beam’ regions spaced over a significant portion of the sky ($\sim 2800 \text{ deg}^2$).

The SDSS Data Release 9 provides estimates of T_{eff} , $\log g$, $[\text{Fe}/\text{H}]$, and $[\alpha/\text{Fe}]$ from an updated and improved version of the SEGUE Stellar Parameter Pipeline (e.g. Lee et al. 2008). Typical uncertainties in the derived parameters are $\sim 4 \text{ km s}^{-1}$ for line-of-sight velocities, and $\sim 0.2 \text{ dex}$ in $[\text{Fe}/\text{H}]$ metallicity.

3 TRIAND AND A SKINNY FRIEND IN THE SPLASH ‘TRASH’

In this section, we analyse the TriAnd overdensity using the MW foreground contamination from the SPLASH survey.

In Fig. 1, we show the spectroscopic footprint of the SPLASH survey fields used in this work. The open squares highlight fields that targeted M31 stars using Washington photometry filters M , T_2 , and $DD051$ to distinguish between dwarfs and giants (see Majewski et al. 2000). Additional fields that used CFHT/MegaCam or Subaru/Suprime-Cam images to select targets are shown with the line-filled squares and tilted filled rectangles, respectively. The location of the recently discovered PAndAS stream is shown in green (Martin et al. 2014). Several of the SPLASH fields coincide with this structure (highlighted in red), and we explore these fields in more detail below.

Fig. 2 shows the V , I colour–magnitude diagram (CMD) for stars in the M31 fields with available Washington photometry. The Washington filters have been converted to Johnson–Cousins bandpasses using the relations given in Majewski et al. (2000). For the SPLASH survey, M31 giant stars were given highest priority on the DEIMOS

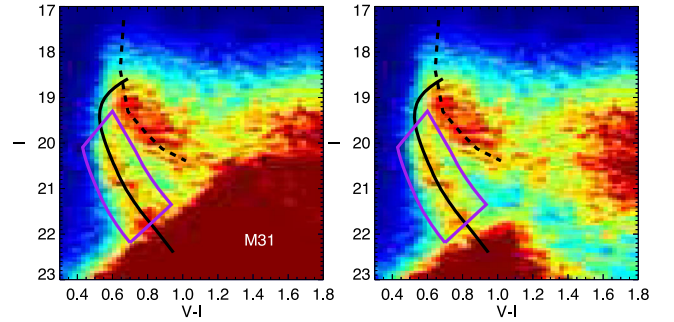


Figure 2. CMD for stars in the M31 fields. All stars are shown in the left-hand panel. Stars with a relatively low probability of being a giant ($g_{\text{prob}} < 0.5$; see Ostheimer 2003) are shown in the right-hand panel. The dashed line indicates the apparent Monoceros Ring ridgeline given by Ibata et al. (2003). We note, however, that this structure could also belong to the (thick) disc population. The solid line shows a Padova isochrone with age $T = 10 \text{ Gyr}$ and metallicity $[\text{Fe}/\text{H}] = -1.5$ shifted to $D = 17 \text{ kpc}$. In the following analysis of the SPLASH data, we select stars within the purple selection box as our foreground MW/TriAnd sample.

masks, but MW foreground targets were also included to fill the mask space. In fact, for the sparsely populated halo fields targeted by Gilbert et al. (2012), the MW foreground stars often outnumber the M31 giants. The right-hand panel excludes high-probability giants from the CMD; bright ($I \sim 20$) MW main-sequence turn-off (MSTO) stars are generally well separated from the M31 giants in CMD space. Two prominent foreground features are visible in the CMD. The feature visible at brighter magnitudes ($I \sim 19$) could be the Monoceros Ring (Newberg et al. 2002; Ibata et al. 2003), or an extension of the thick disc population. For illustration, we indicate the ridgeline of this feature given by Ibata et al. (2003). At fainter magnitudes a second prominent feature is apparent. This is the TriAnd overdensity, which was first noticed in photometric surveys towards M31 by Majewski et al. (2004). The solid line shows a Padova isochrone (Bressan et al. 2012) with age $T = 10 \text{ Gyr}$ and metallicity $[\text{Fe}/\text{H}] = -1.5$ shifted to $D = 17 \text{ kpc}$. This isochrone matches the observed distribution in CMD space very well. In the following analysis of the SPLASH data, we select ($N = 158$) stars within the purple selection box as our foreground MW/TriAnd sample.

3.1 Velocity structure of MW populations

The Galactocentric velocity distributions of stars in the MW MSTO star selection box are shown in the left-hand panels of Fig. 3. Observed heliocentric velocities are converted to Galactocentric ones (V_{GSR}) by assuming a circular speed of 235 km s^{-1} at the position of the Sun ($R_0 = 8.5 \text{ kpc}$) with a solar peculiar motion (U, V, W) = (11.1, 12.24, 7.25) km s^{-1} (Schönrich, Binney & Dehnen 2010). Here, U is directed towards the Galactic centre, V is positive in the direction of Galactic rotation and W is positive towards the North Galactic Pole. The two panels show the velocity distributions ‘on’ and ‘off’ the PAndAS stream. In both panels, there is an excess of stars with velocities at $V_{\text{GSR}} \sim 50 \text{ km s}^{-1}$. We fit a two-component Gaussian profile to the velocity distribution. This includes a fixed MW halo population with $V_{\text{halo}} = 0 \text{ km s}^{-1}$ and $\sigma_{\text{halo}} = 115 \text{ km s}^{-1}$ (e.g. Xue et al. 2008; Brown et al. 2010), and a cold component with mean velocity and dispersion as free parameters. We use a maximum likelihood algorithm to estimate the best-fitting Gaussian parameters for the cold component, and the normalizations for both components are computed iteratively for

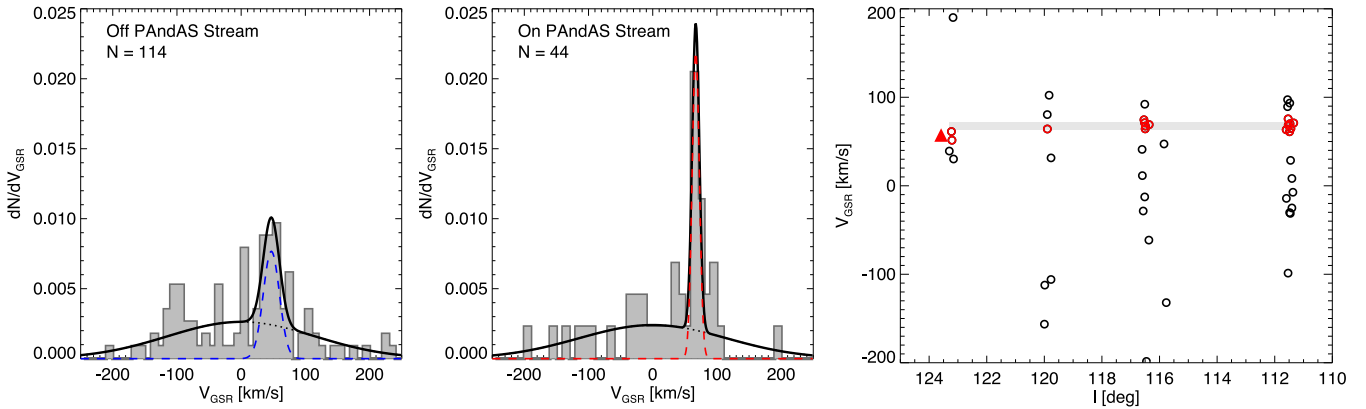


Figure 3. Left and middle panels: velocity distributions of stars in the TriAnd CMD selection box. Stars in fields ‘off’ and ‘on’ the PAndAS stream are shown in the left and middle panels, respectively. The maximum likelihood velocity distributions for MW halo and substructure components are shown with the dotted black and dashed (blue/red) lines. The cold component in the fields not coincident with the PAndAS stream (‘off’ fields) is TriAnd. Right-hand panel: the line-of-sight Galactocentric velocities of stars in the PAndAS stream fields as a function of Galactic longitude. Red circle symbols indicate stars that likely belong to the PAndAS stream ($p_{\text{stream}} > 0.5$). Note that in computing p_{stream} , we assume a constant density of stars with Galactic longitude. The complicated selection function used to target these stars means that it is difficult to draw conclusions from the varying number of stream stars from field to field. The red triangle shows the average velocity for the PAndAS stream found by Martin et al. (2014). There is tentative evidence for a velocity gradient along the PAndAS stream with decreasing l , but the number of stars at high Galactic longitude is very low.

Table 1. The maximum likelihood parameters for the two/three-component (halo/disc plus substructure) Gaussian fits to the velocity distributions. The last column gives the log likelihood ratio for Gaussian distributions with (\mathcal{L}_{SUB}) and without ($\mathcal{L}_{\text{NOSUB}}$) a cold component.

Survey	Substructure	Type	ℓ, b	N_{stars}	$f_{\text{halo}}(f_{\text{disc}})$	$\langle V_{\text{GSR}} \rangle [\text{km s}^{-1}]$	$\sigma [\text{km s}^{-1}]$	$\ln(\mathcal{L}_{\text{SUB}}/\mathcal{L}_{\text{NOSUB}})$
SPLASH	PAndAS stream	MSTO	115.7, −15.9	44	0.69	67^{+8}_{-8}	4^{+20}_{-4}	16.1
SPLASH	TriAnd	MSTO	123.1, −25.7	114	0.76	47^{+6}_{-8}	12^{+13}_{-6}	14.7
SEGUE	TriAnd	BS	149.4, −27.8	78	0.42	36^{+6}_{-6}	21^{+6}_{-5}	29.4
SEGUE	TriAnd	BS	172.3, −33.0	66	0.35	17^{+6}_{-8}	19^{+12}_{-6}	32.7
SEGUE	TriAnd	Giants	150.0, −25.4	287	0.52(0.15)	39^{+14}_{-6}	20^{+14}_{-6}	39.8
SEGUE	TriAnd	MSTO	123.0, −32.5	506	0.80	51^{+7}_{-8}	21^{+10}_{-7}	24.7

each set of parameters ensuring that the total number of objects is correct. The maximum likelihood parameters are listed in Table 1.

The Galactocentric velocity of the structures both on and off the PAndAS stream fields are very similar to the velocity of TriAnd found by Rocha-Pinto et al. (2004) using M-giants (see Fig. 10). The velocity dispersion of the TriAnd feature is hotter than the PAndAS stream. The PAndAS stream has a very narrow velocity dispersion ($\sigma \sim 4 \text{ km s}^{-1}$), in good agreement with the findings of Martin et al. (2014). The proximity of this ‘skinny’ structure to the TriAnd overdensity, in both positional space (i.e. distance and position on the sky) and line-of-sight velocity, strongly suggests that they are associated. Note that a possible association between TriAnd and the PAndAS stream is also mentioned by Sheffield et al. (2014). We discuss the relation between these stellar structures further in Section 5.

The velocities of the PAndAS stream stars as a function of Galactic longitude are shown in the right-hand panel of Fig. 3. Based on the Gaussian velocity distribution for the cold component, we can estimate the probability of stream membership, i.e.

$$p_{\text{sub}} = \frac{f_{\text{sub}} G_{\text{sub}}}{f_{\text{sub}} G_{\text{sub}} + f_{\text{halo}} G_{\text{halo}}}$$

$$G = \frac{1}{\sqrt{2\pi}\sigma} \exp\left(-\frac{(v - \langle v \rangle)^2}{2\sigma^2}\right), \quad (1)$$

where p_{sub} is the probability for the relevant substructure (in this case the stream), G is a Gaussian velocity distribution, and f is the

fractional contribution of the substructure/halo components. The circles highlighted in red show those stars with $p_{\text{stream}} > 0.5$. There is tentative evidence for a velocity gradient along the PAndAS stream with decreasing ℓ , but the number of stars at high Galactic longitude is very low.

3.2 Relative metallicity of MW populations

We now consider if there is any significant metallicity difference between the TriAnd feature and the PAndAS stream. The SPLASH spectra cover the wavelength range of the Ca II triplet ($\lambda 8500 \text{ \AA}$), which was used in previous work to estimate the metallicity of M31 RGB stars (e.g. Gilbert et al. 2006; Guhathakurta et al. 2006; Ho et al. 2014). For the bluer, main-sequence stars considered here, the Ca II lines are much weaker than for redder stars and the signal-to-noise ratio (S/N) of the selected spectra are relatively low ($S/N \sim 5 \text{ \AA}^{-1}$). Thus, in order to maximize any apparent metallicity signal in the data, we stack together spectra in $V - I$ colour-bins for each of the different MW substructures. We consider the high-probability members of the TriAnd overdensity (with $P_{\text{triand}} > 0.5$), PAndAS stream ($P_{\text{stream}} > 0.5$) and field halo ($P_{\text{halo}} > 0.5$). This corresponds to $N = 36, 15$, and 107 objects for each of these populations, respectively.

Fig. 4 shows the effective Ca II equivalent width (EW) for stacked SPLASH spectra in the field halo, the PAndAS stream and TriAnd. Only two of the Ca II triplet lines are used as the line at $\lambda 8498 \text{ \AA}$ is very weak in the low S/N, hot spectra of the turn-off stars.

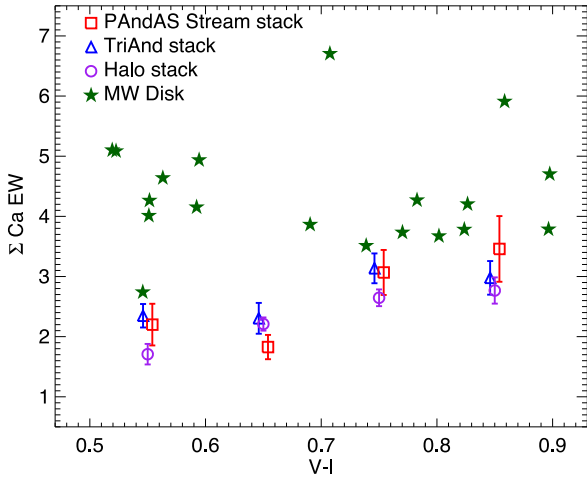


Figure 4. Calcium triplet EWs as a function of $V - I$ colour. The green stars are high-S/N MW disc stars. The stacked values for the PAndAS stream, TriAnd, and field MW halo stars are shown with the red squares, blue triangles, and purple circles, respectively. The bins in $V - I$ colour are 0.1 dex wide, and the central values for the of the bin are shown for the x-axis values. The symbols are shown slightly offset on the x-axis for clarity. The average EWs for the stack, and associated error in the mean, are shown with the y-axis values. We find no significant difference in metallicity between the PAndAS stream, TriAnd and field halo stars based on the weak Ca II lines.

In addition, it occasionally landed in the inter-CCD gap of the DEIMOS detector for some of the observing runs (Gilbert et al. 2006). The EW of the Ca II (Σ Ca) triplet was calculated using a linear combination of the EWs of the remaining two lines that has been shown to maximize the S/N of the feature (Rutledge, Hesser & Stetson 1997): $\Sigma\text{Ca} = 1.0\text{EW}(\lambda 8542 \text{ \AA}) + 0.6\text{EW}(\lambda 8662 \text{ \AA})$. We note that we use the Ca EW as a proxy for metallicity, but do not enforce a calibration between Ca EW and $[\text{Fe}/\text{H}]$ as this relationship has (to our knowledge) only been optimized for giant stars.

We find no significant difference between the three populations. For comparison, high-S/N spectra for (likely) bright MW disc stars (selected from the same SPLASH fields with $I < 19.5$, $V_{\text{GSR}} > 100 \text{ km s}^{-1}$) are shown with the green star symbols. These stars, which are presumably more metal rich, have systematically higher Ca II EWs than the halo populations. This suggests that the weak Ca II lines are able to differentiate between significant metallicity differences for blue ($V - I < 1.0$) stars. However, it is unclear whether these weak Ca II lines are able to distinguish between modest metallicity differences for the (metal-poorer) halo populations. We defer a more detailed study of the metallicities of the TriAnd and PAndAS stream structures in the SPLASH survey to future work.

4 TRIAND IN SDSS/SEGUE: PROXIMITY TO THE SEGUE 2 SATELLITE

In the *Segue* 2 discovery paper, Belokurov et al. (2009) noted the proximity of this satellite to TriAnd. Inspired by this possible connection, we now investigate the TriAnd overdensity using A-type, giant and MSTO stars in the vicinity of the *Segue* 2 satellite.

In the left-hand panel of Fig. 5, we show the (large-scale) distribution of M-giants selected from *WISE* and 2MASS photometry (Koposov et al., in preparation). M-giants in the approximate distance range $15 \lesssim D/\text{kpc} \lesssim 30$, are selected with the following

colour and magnitude cuts:

$$\begin{aligned} -0.2 < W_1 - W_2 < -0.07 \\ 0.9 < J - K_s < 1.3 \\ 11 < K_s < 12. \end{aligned} \quad (2)$$

The distribution of these M-giants in equatorial coordinates is shown in Fig. 5. The excess of M-giants near the TriAnd region (RA, Dec. $\sim 30, 30$ deg) is the TriAnd overdensity that Rocha-Pinto et al. (2004) discovered using a similar selection of M-giants from 2MASS. The orange star indicates the location of the *Segue* 2 satellite, and the black box indicates the region we explore in the following subsections using SDSS/SEGUE spectra.

4.1 A-type stars

A-type stars consist mainly of blue horizontal branch (BHB) and blue straggler (BS) stars located bluewards of the MSTO. These stars are useful halo tracers owing to the limited contamination by other stellar populations (cf. dwarfs and giants at redder colours) and their relatively bright absolute magnitudes. We selected A-type stars from the SEGUE spectroscopic sample by applying the following colour and magnitude cuts:

$$\begin{aligned} -0.5 < g - r < 0.1 \\ 0.5 < u - g < 2.0 \\ 18 < g < 20. \end{aligned} \quad (3)$$

BS stars dominate over BHB stars in this faint magnitude slice (Deason et al. 2014), and although they are ~ 2 mag fainter than BHB stars, BS stars have proved a useful tool to study the prominent Sgr stream in the MW halo (Belokurov et al. 2014). The magnitude cut is applied to select BS stars with approximate distances of $10 \lesssim D/\text{kpc} \lesssim 30$ (Deason, Belokurov & Evans 2011), roughly coincident with the distance of the TriAnd overdensity (Majewski et al. 2004; Rocha-Pinto et al. 2004; Sheffield et al. 2014).

In the right-hand panel of Fig. 5, we show the spectroscopic SEGUE fields in the vicinity of the *Segue* 2 satellite and M31. The coloured points show the A-type stars selected with the above colour cuts. We only consider stars with $|b| > 20$ deg to avoid significant extinction corrections and disc contamination. The location of the *Segue* 2 satellite is shown with the orange star symbol. We also indicate the location of M31 and the PAndAS stream. The dashed line illustrates a great circle through the PAndAS stream, which coincides with some of the SEGUE fields.

The Galactocentric velocity distribution of A-type stars in the SEGUE fields are shown in Fig. 6. The location of the fields are colour-coded to match Fig. 5. Significant velocity signatures coincident with the TriAnd overdensity are seen in the fields closest to the *Segue* 2 satellite, and roughly coincident with the great circle through the PAndAS stream. The fields at lower RA have less significant signatures; however, there is a hint of the TriAnd velocity feature in the top-right panel. Note that in Section 4.4, we show that a more prominent signature of the TriAnd overdensity is visible for MSTO stars in this region of the sky (see the middle-left panels of Fig. 9). We perform a two-component (halo plus sub-structure) Gaussian fit to the velocity distributions in the fields near *Segue* 2 using the same maximum likelihood procedure as outlined in Section 3.1. The double-Gaussian fit is shown in Fig. 6 and the maximum likelihood parameters are given in Table 1.

As noted previously by Belokurov et al. (2009), the *Segue* 2 satellite has a very similar systemic velocity to TriAnd and is located at a comparable distance ($D \sim 30$ kpc). This, in addition to the

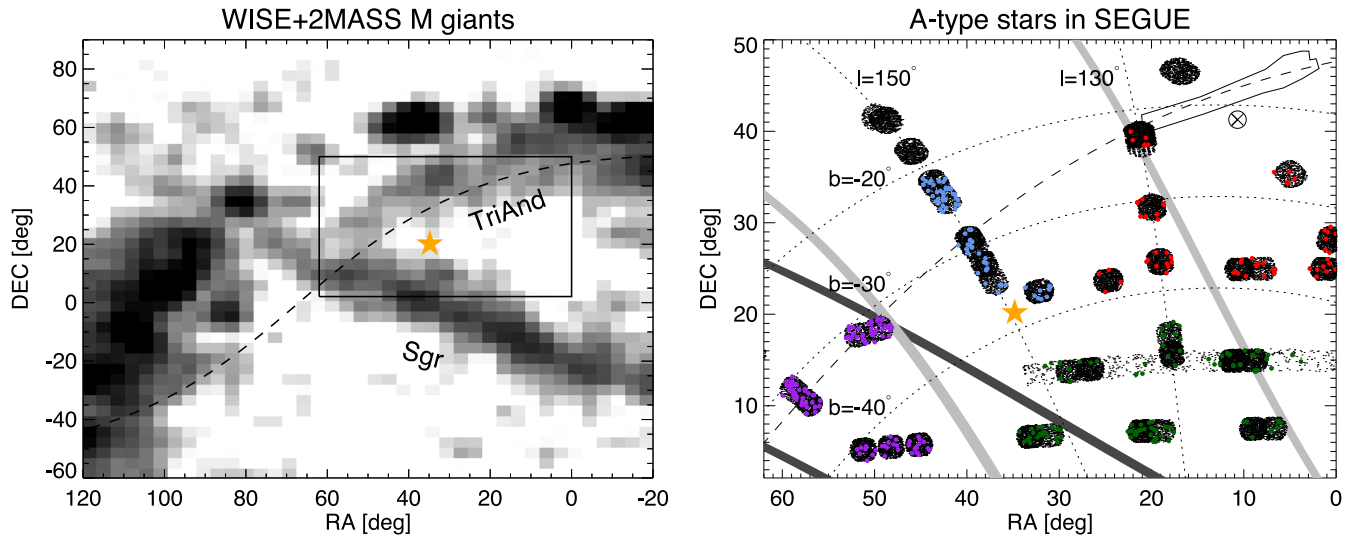


Figure 5. Left-hand panel: density of M-giants selected from *WISE* and 2MASS photometry (Koposov et al., in preparation) in equatorial coordinates. The black box indicates the region where we analyse TriAnd in the SDSS/SEGUE spectroscopic fields. The dashed line illustrates a great circle through the PAndAS stream. Right-hand panel: the SEGUE spectroscopic footprint in equatorial coordinates. Fields close to the *Segue 2* satellite are shown, and the location of the satellite is indicated by the orange star symbol. A-type stars (mainly BSs) selected in these fields are highlighted in blue, red, purple, and green according to their equatorial position on the sky. The location of the PAndAS stream and the centre of M31 are highlighted in the top-right corner. Lines of constant Galactic latitude/longitude are shown with dotted lines. The approximate tracks of Sgr and Cetus are shown with the dark and light grey shaded regions, respectively (Belokurov et al. 2014).

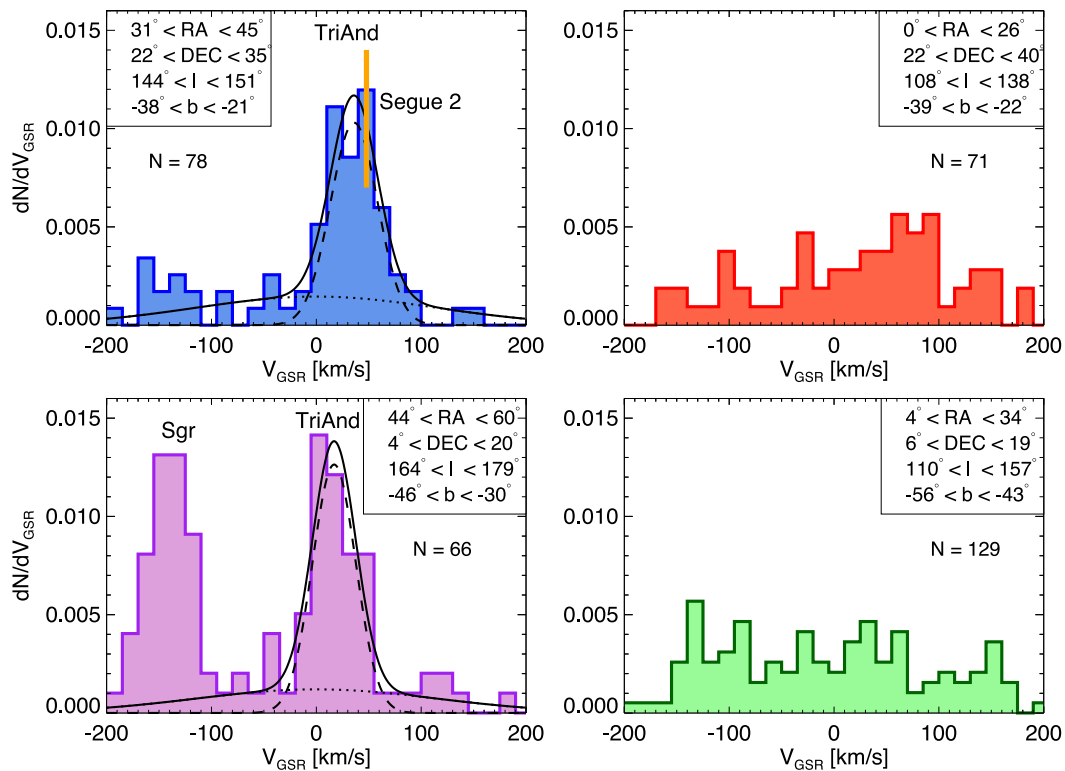


Figure 6. The Galactocentric velocity distributions of A-type stars located in the selected SEGUE fields. The fields coincident with a great circle through the PAndAS stream, show a significant excess at $V_{\text{GSR}} \sim 30 \text{ km s}^{-1}$, which is coincident with the velocity of the TriAnd overdensity and the *Segue 2* satellite. The solid line shows a two-component Gaussian fit to the velocity distribution, where the field halo velocity distribution (dotted line) is fixed with $V_{\text{halo}} = 0 \text{ km s}^{-1}$ and $\sigma_{\text{halo}} = 115 \text{ km s}^{-1}$. The TriAnd overdensity is less prominent in the fields at lower RA. Velocity signatures associated with Sgr are also highlighted. The velocity distributions of TriAnd and Sgr in the bottom-left panel have similar dispersions ($\sim 20 \text{ km s}^{-1}$), suggesting that these may be similar mass structures.

PAndAS stream, is the second stream/satellite that has been found in close proximity (both spatially, and in velocity space) to the large TriAnd overdensity. It is plausible that these structures are all associated, and we discuss this further in Section 5.

4.2 Giant stars

Giants are useful tracers of halo populations owing to their bright absolute magnitudes, which allows relatively shallow spectroscopic surveys to probe out to large distances in the halo. Giant stars in SEGUE are selected using a similar selection to Belokurov et al. (2014), but we include a wider colour range to include lower $[\text{Fe}/\text{H}]$ stars:

$$\begin{aligned} 17 < g < 20 \\ 0.6 < g - i < 2.0 \\ 3.66 < \log(T_{\text{eff}}/\text{K}) < 3.76 \\ 1.5 < \log(g_s) < 3.8. \end{aligned} \quad (4)$$

We focus on the SEGUE stripe with approximately constant $\ell = 150$ deg (see Fig. 5), which contains the most obvious TriAnd signal and coincides with the *Segue 2* satellite. In Fig. 7, we show the velocity distribution of these giants as a function of Galactic latitude (top panel) and $[\text{Fe}/\text{H}]$ metallicity (bottom panel). The black curve across the top panel is the predicted thick disc signature for dwarf contaminants in the sample. At higher latitudes, the thick disc has similar velocity to the TriAnd structure so it is difficult to distinguish these components. The Sgr stream is prominent at latitudes away from the Galactic plane and negative Galactocentric velocities. In the bottom panel, the velocity distribution is shown as a function of metallicity. In this panel, only stars with latitudes $b > -40$ deg (and $b < -13$ deg) are shown to limit confusion with the thick disc and Sgr. The velocity of the *Segue 2* satellite is indicated and, as mentioned above, is very similar to the TriAnd overdensity. The metallicity of TriAnd is more metal rich than the field halo stars, but it has a very broad distribution with a tail to lower metallicities. In contrast, the thick disc distribution is clustered around relatively high metallicities ($[\text{Fe}/\text{H}] \sim -0.5$) with no significant low-metallicity tail.

4.3 Metallicity and distances of A-type stars and giants

We study the metallicity distributions of the halo populations further in Fig. 8. In the left-hand panel, the grey histogram shows the metallicity distribution of A-type stars selected in SEGUE with $20 < |b|/\text{deg} < 40$. High-probability TriAnd stars in the $\ell \sim 150$ deg SEGUE fields ($P_{\text{triand}} > 0.5$) are shown with the line-filled light blue histogram. The TriAnd stars seem to be biased to higher metallicities. It is worth noting that the metallicity estimates for A-type stars likely suffer from significant systematic biases, and the $[\text{Fe}/\text{H}]$ estimates from the SEGUE pipeline were not designed for these hot, A-type stars. None the less, these metallicities should be sufficient to look at relative differences between populations, and indeed, we find that the stars belonging to TriAnd tend to be more metal rich than the field halo.

The navy blue dashed line shows the metallicity distribution of giant stars (with $\ell \sim 150$ deg) in SEGUE with high probability of belonging to TriAnd. Encouragingly, these metallicities agree very well with the A-type stars in TriAnd. We note that Rocha-Pinto et al. (2004) and Sheffield et al. (2014) found that their samples of M-giants in TriAnd typically have metallicities of $[\text{Fe}/\text{H}] \sim -1.0$, in good agreement with our findings.

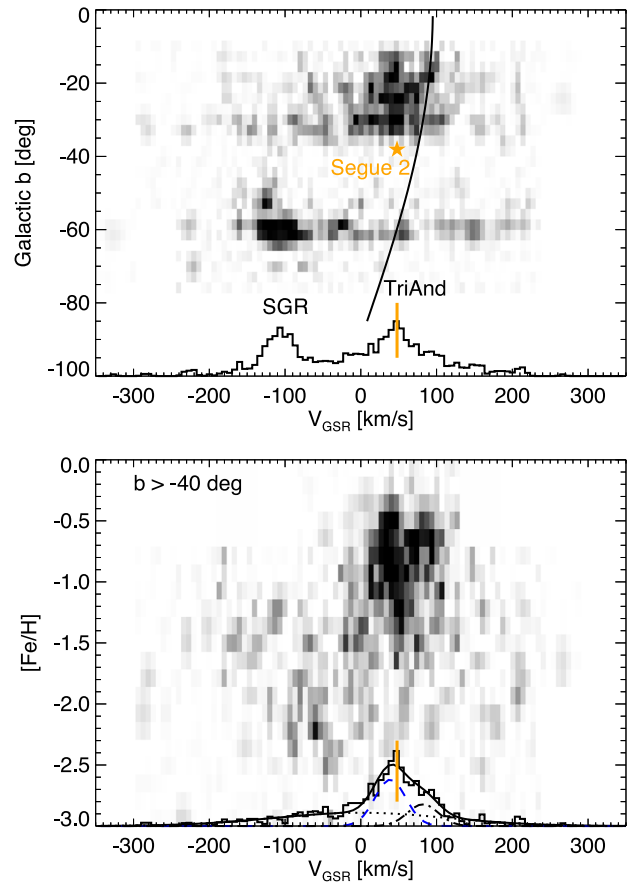


Figure 7. The Galactocentric velocity distribution of giant stars selected from SEGUE with Galactic longitude $\ell \sim 150$ deg. Top panel: the velocity distribution at different Galactic latitudes. The Sgr dwarf is prominent at latitudes away from the Galactic plane ($b < -40$ deg) and the TriAnd overdensity is significant at latitudes closer to the Galactic plane ($b > -40$ deg). The orange star symbol indicates the Galactocentric velocity and Galactic latitude of the *Segue 2* satellite. The solid black line shows the velocity track of the thick disc as a function of Galactic latitude. Bottom panel: the velocity distribution of SEGUE giants as a function of $[\text{Fe}/\text{H}]$ metallicity. Only stars with $b > -40$ deg (and $b < -13$ deg) are shown to limit contamination from Sgr and thick disc stars with a similar velocity to TriAnd. The TriAnd feature is more metal rich than the field halo giants, but has a wide distribution of metallicities. The solid line shows a triple Gaussian fit to the velocity distribution, which includes contributions from the halo, thick disc and TriAnd. The Gaussian parameters for the halo and thick disc are kept fixed in the fitting procedure. The blue dashed line shows the contribution from TriAnd.

In the right-hand panel of Fig. 8, we show the approximate distance distributions of the SEGUE stars likely related to TriAnd. For the A-type stars, we use the Deason et al. (2011) absolute magnitude calibration for BS stars assuming a metallicity of $[\text{Fe}/\text{H}] \sim -1.0$. The distances of the BS giant stars are shown by the light blue, line-filled region. We note that a small fraction of these stars may be BHBs, and these will have distances $D \gtrsim 30$ kpc. To estimate the distances of the giant stars, we cross match these stars with the SEGUE K-giant sample compiled by Xue et al. (2014), who provide distance estimates to the giants. Only 6 (out of 111) stars are in the Xue et al. (2014) sample and their approximate distances are shown with the navy blue dashed line in the figure. Note that the small fraction of overlap between these samples is likely because

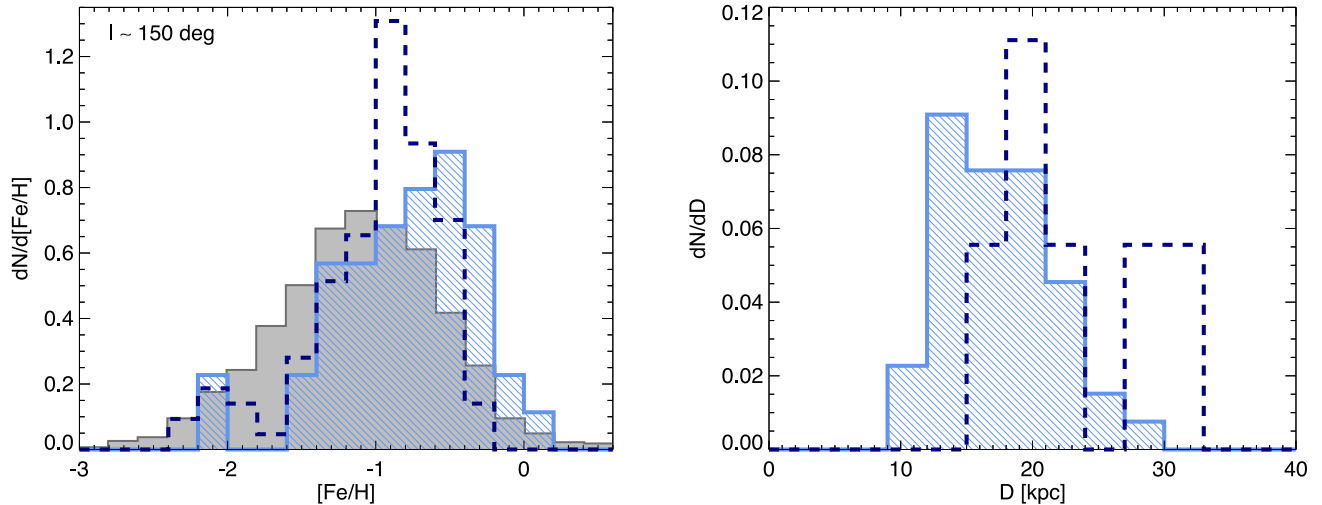


Figure 8. Left-hand panel: the metallicity distribution of ($N = 2606$) A-type stars in SEGUE with $18 < g < 20$ and $20 < |b|/\text{deg} < 40$ (grey filled histogram). The distribution for ($N = 44$) SEGUE A-type stars in the TriAnd overdensity are shown with the light blue line-filled histogram. TriAnd is clearly more metal rich than the overall halo population. The navy blue dashed line shows the metallicity distribution for ($N = 107$) SEGUE giant stars in the TriAnd overdensity. The stars (both A-type and giants) in the TriAnd overdensity are selected from the $\ell \sim 150$ deg fields (see Table 1) with $p_{\text{triand}} > 0.5$. Right-hand panel: the approximate distances for the BS stars and giants in the TriAnd overdensity near $\ell \sim 150$ deg. The absolute magnitude calibration from Deason et al. (2011) is used for the ($N = 44$) BS stars, assuming a metallicity of $[Fe/H] \sim -1.0$. We use the Xue et al. (2014) distance estimates for the giant stars; only ($N = 6$) SEGUE giants which overlap with this catalogue are shown in the figure.

the Xue et al. (2014) K-giants are biased towards lower metallicity giants (e.g. less than 10 per cent of the sample have $[Fe/H] > -1.0$).

For both populations, BSs and giants, the distance estimates are highly uncertain, and we caution that these are only approximate values. However, it is encouraging that the BS stars and giants are in good agreement, with distances of $D \sim 20$ kpc, very similar to previous estimates for the TriAnd overdensity (Majewski et al. 2004; Rocha-Pinto et al. 2004; Sheffield et al. 2014). We also note that our selection of A-type stars and giants identifies signatures of the Sgr stream, located at $D \sim 30$ kpc in this region of the sky, therefore giving further support to our distance calibration.

4.4 Main sequence turn-off stars

We end this section by looking at MSTO stars in the SDSS/SEGUE fields. This brings us right back to where we started in Section 3, where we used MSTO stars in the SPLASH survey to probe the TriAnd overdensity. MSTO stars are fainter than A-type stars and giants, but there are many more of them. In fact, the high number density of MSTO stars has been exploited in previous work to map out streams and structures in the MW halo over large regions of the sky (e.g. Belokurov et al. 2006a).

In the top panels of Fig. 9, we show CMDs for stars in SDSS imaging. Each panel shows the different regions of the sky that we probed in the previous subsections (see Fig. 5). In the fields closer to the Galactic plane ($b \sim -30$ deg) the TriAnd MSTO is clear. A well-defined MS for TriAnd is also apparent, which appears to be limited in distance as opposed to the blurry mess of halo stars spanning a large range of distances. In the left-hand panel ($(\ell, b) \sim (149, -30)$ deg), a gap between the more distant TriAnd MSTO and the disc population is evident. Furthermore, the TriAnd turn-off is slightly bluer, suggesting a metal-poorer population than the disc.

The bottom panels of Fig. 9 show the corresponding velocity distributions for the MSTO stars in each field (selected with $0.1 < g - r < 0.4$). This allows us to distinguish between the

brighter disc stars and the fainter stars in the MSTO of TriAnd. The velocity signatures of the thick disc and TriAnd are apparent in the two left-hand panels. At brighter magnitudes, $16 \lesssim r \lesssim 19$, there is a clear thick disc MSTO signature that has a gentle slope with V_{GSR} due to the projection of progressively more distant stars on to the line of sight. For $r \gtrsim 19$, the TriAnd signature appears clearly disconnected from the thick disc sequence with $V_{\text{GSR}} \sim 40 \text{ km s}^{-1}$. At Galactic longitudes close to $\ell \sim 180$ deg (third panel), the disc and halo signatures are difficult to disentangle as the disc crosses $V_{\text{GSR}} = 0 \text{ km s}^{-1}$. Finally, the fourth panel displays a much broader halo population at fainter magnitudes ($\sigma \sim 100 \text{ km s}^{-1}$), in good agreement with Fig. 6.

For the MSTO stars located in the second panel (indicated in red, $(\ell, b) \sim (123, -32)$ deg), we perform a two-component (halo plus substructure) fit to the velocity distribution of MSTO stars with $19 < r < 20$. The magnitude cut is included to remove the majority of disc stars. The resulting Gaussian parameters for the TriAnd feature are given in Table 1. The results are in good agreement with our findings from the SPLASH survey (see Section 3.1), which is located in a similar region on the sky (see red points in Fig. 5).

5 A FOSSIL RECORD OF GROUP INFALL?

The results outlined in the previous sections suggest that at least three distinct substructures in the MW are associated, namely the large TriAnd overdensity, the PAndAS stream and the *Segue 2* satellite.

In Fig. 10, we give a summary of our results, showing Galactocentric velocity as a function of Galactic longitude for the various halo populations studied in this work. The shaded regions indicate the 1σ velocity distributions about the mean (see Table 1) and span the extent in Galactic longitude of each field. Also shown are the *Segue 2* satellite, the velocity measurement for the PAndAS stream found in Martin et al. (2014), and the TriAnd M-giant samples compiled by Rocha-Pinto et al. (2004) and Sheffield et al. (2014). Note that the velocity dispersion we find for the TriAnd features

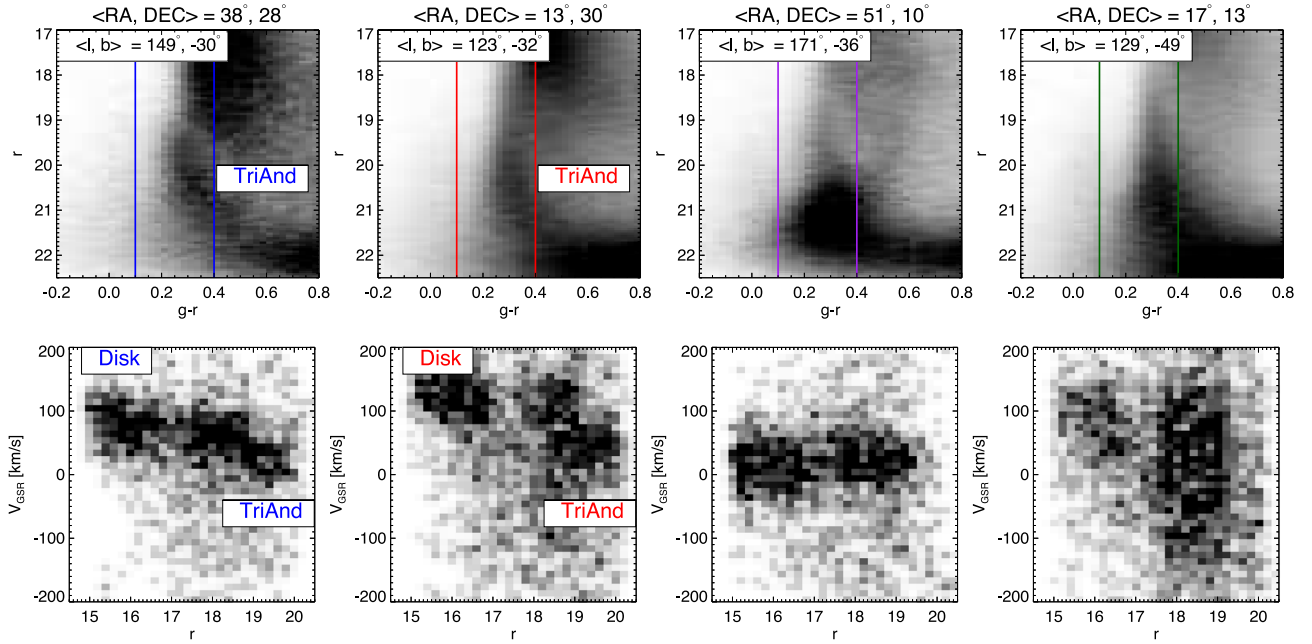


Figure 9. Top panels: CMDs for stars in SDSS. Each panel is for different regions on the sky (average equatorial and Galactic coordinates are indicated) with the same colour coding as Fig. 5. The solid lines show the $g - r$ colour region used to select MSTO stars. The TriAnd MS is clear in the left-hand panels (cf. Fig. 2). The MSTO of the Sgr is evident in the third panel at $r \sim 21$. Bottom panels: Galactocentric velocity against (r -band) magnitude for MSTO stars. Each panel is for different regions on the sky as indicated in the top panels. The TriAnd overdensity at $V_{\text{GSR}} \sim 40 \text{ km s}^{-1}$ is clearly visible at $r \sim 19\text{--}20$. This velocity signature is distinct from the brighter disc stars.

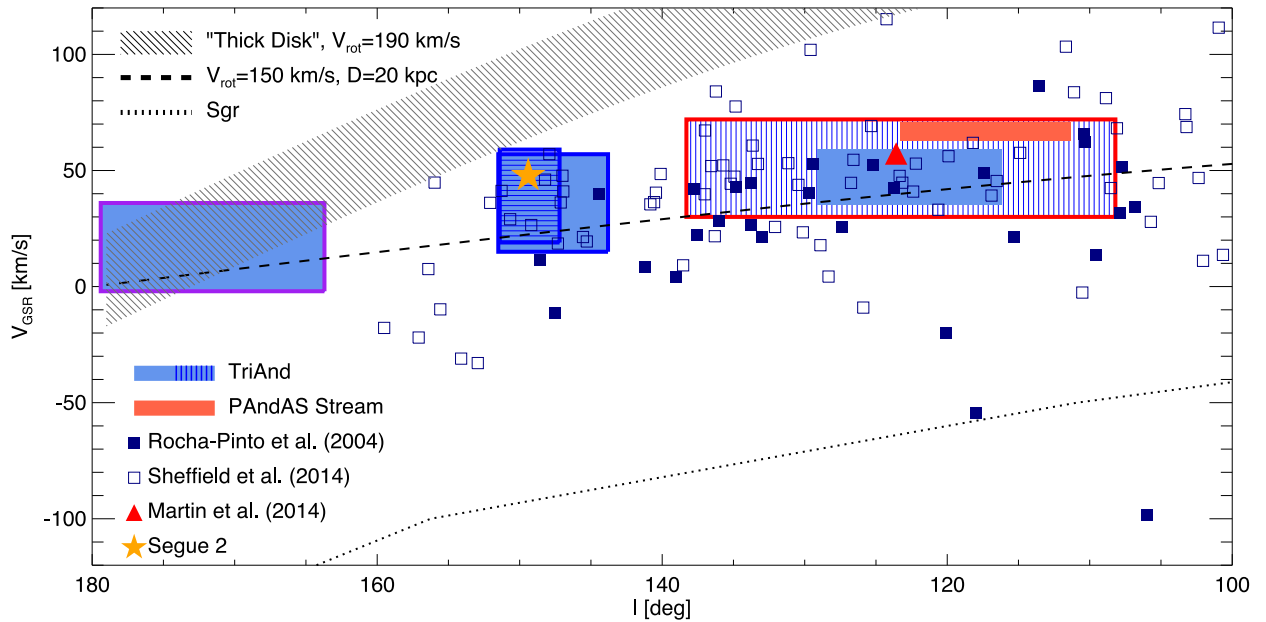


Figure 10. Galactocentric velocity versus Galactic longitude for TriAnd and its ‘siblings’. The blue shaded (solid and line-filled) regions show the velocities of TriAnd in the SPLASH and SEGUE fields, and the red shaded region indicates the velocity of the PAndAS stream. The locations of the TriAnd fields for A-type stars in SEGUE are colour-coded (blue, purple solid lines) as in Fig. 5. These shaded regions indicate the 1σ velocity distributions about the mean, and span the extent in Galactic longitude of each field. The horizontal line-filled region near the *Segue 2* satellite shows the velocity distribution of TriAnd traced by giant stars. The vertical line-filled region at lower Galactic longitude shows the velocity distribution of TriAnd traced by MSTO stars. The red outline follows the same colour-coding as Fig. 5 to indicate the location of the field. The dark blue squares show the velocities of the TriAnd M-giants compiled by Rocha-Pinto et al. (2004, filled) and Sheffield et al. (2014, open). The average velocity found by Martin et al. (2014) for the PAndAS stream is shown with the red triangle symbol, and the *Segue 2* satellite (Belokurov et al. 2009) is shown with the orange star symbol. The dashed line indicates a circular orbit with $V_{\text{rot}} = 150 \text{ km s}^{-1}$, and the line-filled grey region shows a local thick disc-like population. The approximate projection of the Sgr stream orbit on to V_{GSR} , ℓ space is indicated with the dotted line (Belokurov et al. 2014).

($\sigma \sim 20 \text{ km s}^{-1}$) is similar to the spread in velocities of the M-giant samples. For illustration, the dashed line shows the Galactocentric velocity profile for planar rotation at $D = 20 \text{ kpc}$ (corresponding to Galactocentric distance, $r \sim 25 \text{ kpc}$) with $V_{\text{rot}} = 150 \text{ km s}^{-1}$, and the line-filled region shows the approximate profile for a local thick disc population. The approximate projection of the Sgr stream orbit on to V_{GSR}, ℓ space is also indicated with the dotted line (Belokurov et al. 2014).

It is clear that TriAnd, the PAndAS stream, and *Segue 2* may all follow very similar orbits, possibly even forming a ‘ring-like’ structure with a roughly planar orbit ($V_{\text{rot}} \approx V_{\phi} \approx 150 \text{ km s}^{-1}$), although without proper motion information we are unable to rule out different orbital configurations.² This suggests that these substructures fell into the MW halo as a group, with the large, metal-rich TriAnd structure dominating the group mass (i.e. the central) and the PAndAS stream and *Segue 2* being satellites of TriAnd.

Wetzel, Deason & Garrison Kimmel (in preparation) recently showed that a significant fraction (~ 30 per cent) of ultrafaint dwarfs likely fell into the MW as satellites of larger dwarfs. This work also shows that group infall is more common for lower mass subhaloes. Thus, it is perhaps unsurprising that the ‘least massive galaxy’, *Segue 2*, can be likened to a sub-subhalo in the hierarchical structure formation paradigm.

If *Segue 2* was indeed born as a subsystem of a larger dwarf, then it may have undergone pre-processing by its original host’s tidal field. This may explain why *Segue 2* lies significantly off the metallicity–stellar mass relation that is closely followed by most MW dwarfs (Kirby et al. 2013b). However, it is also worth noting that, even if *Segue 2* has no dynamical association with TriAnd, there is likely significant contamination from this large, metal-rich structure along the line of sight towards *Segue 2*. This could affect both the velocity dispersion measurements of the satellite, and the metallicity distribution of the apparent *Segue 2* members, if not properly accounted for.

6 CONCLUSIONS

We have investigated the large, diffuse TriAnd overdensity using foreground MSTO star contamination in the SPLASH survey of M31, and using A-type stars, giants and MSTO stars from the SEGUE survey in the vicinity of the *Segue 2* satellite. Our main conclusions are summarized as follows.

(i) Several of the SPLASH spectroscopic fields coincide with the recently discovered PAndAS stream, which is located at a similar distance ($D \sim 20 \text{ kpc}$) and region of the sky to TriAnd. We find that the velocity of the stream is very similar to TriAnd ($V_{\text{GSR}} \sim 50 \text{ km s}^{-1}$), and there is tentative evidence for a velocity gradient in Galactic longitude along the PAndAS stream that is similar to the apparent gradient of the TriAnd overdensity over a much larger area of the sky.

(ii) At larger Galactic longitudes ($\ell \sim 150 \text{ deg}$), the SEGUE fields allow us to probe the TriAnd overdensity in the vicinity of the *Segue 2* satellite, which has previously been claimed to be coincident with this large substructure (Belokurov et al. 2009). Using A-type, giant and MSTO stars, we find that the velocity of TriAnd stars in the vicinity of *Segue 2* are very similar to the systemic velocity of the satellite.

(iii) The metallicity of TriAnd stars in the SEGUE survey are more metal rich than the field halo. The Ca II triplet lines from the SPLASH data have similar EWs in the PAndAS stream, TriAnd and the field halo, in apparent discord with our findings from SEGUE. However, these weak lines may be unable to distinguish between modest metallicity differences, especially for the relatively blue main-sequence stars under consideration.

(iv) We combine the results from the SPLASH and SEGUE surveys to map the TriAnd overdensity over $\sim 80 \text{ deg}$ in Galactic longitude. The velocity profile of the overdensity resembles a ‘ring-like’ structure at $D \sim 20\text{--}30 \text{ kpc}$. The coincidence of the PAndAS stream and *Segue 2* satellite in positional and velocity space to TriAnd strongly suggests that they are all associated, and we advocate that these substructures may be a fossil record of group infall on to the MW halo. In this scenario, the large, metal-rich TriAnd structure is the group central, whilst the PAndAS stream and *Segue 2* are satellites of a MW satellite.

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² The orbit is likely slightly inclined as the velocity signature does not seem to go through $V_{\text{GSR}} = 0 \text{ km s}^{-1}$ at $\ell = 180 \text{ deg}$.

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